



HAL
open science

GW150914: Implications for the stochastic gravitational wave background from binary black holes

B.P. Abbott, G. Cagnoli, J. Degallaix, V. Dolique, R. Flaminio, M. Granata, D. Hofman, C. Michel, R. Pedurand, L. Pinard, et al.

► **To cite this version:**

B.P. Abbott, G. Cagnoli, J. Degallaix, V. Dolique, R. Flaminio, et al.. GW150914: Implications for the stochastic gravitational wave background from binary black holes. *Physical Review Letters*, 2016, 116 (13), pp.131102. 10.1103/PhysRevLett.116.131102 . in2p3-01273211

HAL Id: in2p3-01273211

<https://in2p3.hal.science/in2p3-01273211v1>

Submitted on 10 Jan 2024

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.



CHORUS

This is the accepted manuscript made available via CHORUS. The article has been published as:

GW150914: Implications for the Stochastic Gravitational-Wave Background from Binary Black Holes

B. P. Abbott *et al.* (LIGO Scientific Collaboration and Virgo Collaboration)

Phys. Rev. Lett. **116**, 131102 — Published 31 March 2016

DOI: [10.1103/PhysRevLett.116.131102](https://doi.org/10.1103/PhysRevLett.116.131102)

GW150914: Implications for the stochastic gravitational-wave background from binary black holes

B. P. Abbott,¹ R. Abbott,¹ T. D. Abbott,² M. R. Abernathy,¹ F. Acernese,^{3,4} K. Ackley,⁵ C. Adams,⁶ T. Adams,⁷ P. Addesso,³ R. X. Adhikari,¹ V. B. Adya,⁸ C. Affeldt,⁸ M. Agathos,⁹ K. Agatsuma,⁹ N. Aggarwal,¹⁰ O. D. Aguiar,¹¹ L. Aiello,^{12,13} A. Ain,¹⁴ P. Ajith,¹⁵ B. Allen,^{8,16,17} A. Allocca,^{18,19} P. A. Altin,²⁰ S. B. Anderson,¹ W. G. Anderson,¹⁶ K. Arai,¹ M. C. Araya,¹ C. C. Arceneaux,²¹ J. S. Areeda,²² N. Arnaud,²³ K. G. Arun,²⁴ S. Ascenzi,^{25,13} G. Ashton,²⁶ M. Ast,²⁷ S. M. Aston,⁶ P. Astone,²⁸ P. Aufmuth,⁸ C. Aulbert,⁸ S. Babak,²⁹ P. Bacon,³⁰ M. K. M. Bader,⁹ P. T. Baker,³¹ F. Baldaccini,^{32,33} G. Ballardin,³⁴ S. W. Ballmer,³⁵ J. C. Barayoga,¹ S. E. Barclay,³⁶ B. C. Barish,¹ D. Barker,³⁷ F. Barone,^{3,4} B. Barr,³⁶ L. Barsotti,¹⁰ M. Barsuglia,³⁰ D. Barta,³⁸ J. Bartlett,³⁷ I. Bartos,³⁹ R. Bassiri,⁴⁰ A. Basti,^{18,19} J. C. Batch,³⁷ C. Baune,⁸ V. Bavigadda,³⁴ M. Bazzan,^{41,42} B. Behnke,²⁹ M. Bejger,⁴³ A. S. Bell,³⁶ C. J. Bell,³⁶ B. K. Berger,¹ J. Bergman,³⁷ G. Bergmann,⁸ C. P. L. Berry,⁴⁵ D. Bersanetti,^{46,47} A. Bertolini,⁹ J. Betzwieser,⁶ S. Bhagwat,³⁵ R. Bhandare,⁴⁸ I. A. Bilenko,⁴⁹ G. Billingsley,¹ J. Birch,⁶ R. Birney,⁵⁰ S. Biscans,¹⁰ A. Bisht,^{8,17} M. Bitossi,³⁴ C. Biwer,³⁵ M. A. Bizouard,²³ J. K. Blackburn,¹ C. D. Blair,⁵¹ D. G. Blair,⁵¹ R. M. Blair,³⁷ S. Bloemen,⁵² O. Bock,⁸ T. P. Bodiya,¹⁰ M. Boer,⁵³ G. Bogaert,⁵³ C. Bogan,⁸ A. Bohe,²⁹ P. Bojtos,⁵⁴ C. Bond,⁴⁵ F. Bondu,⁵⁵ R. Bonnand,⁷ B. A. Boom,⁹ R. Bork,¹ V. Boschi,^{18,19} S. Bose,^{56,14} Y. Bouffanais,³⁰ A. Bozzi,³⁴ C. Bradaschia,¹⁹ P. R. Brady,¹⁶ V. B. Braginsky,⁴⁹ M. Branchesi,^{57,58} J. E. Brau,⁵⁹ T. Briant,⁶⁰ A. Brillet,⁵³ M. Brinkmann,⁸ V. Brisson,²³ P. Brockill,¹⁶ A. F. Brooks,¹ D. D. Brown,⁴⁵ N. M. Brown,¹⁰ C. C. Buchanan,² A. Buikema,¹⁰ T. Bulik,⁴⁴ H. J. Bulten,^{61,9} A. Buonanno,^{29,62} D. Buskulic,⁷ C. Buy,³⁰ R. L. Byer,⁴⁰ L. Cadonati,⁶³ G. Cagnoli,^{64,65} C. Cahillane,¹ J. Calderón Bustillo,^{66,63} T. Callister,¹ E. Calloni,^{67,4} J. B. Camp,⁶⁸ K. C. Cannon,⁶⁹ J. Cao,⁷⁰ C. D. Capano,⁸ E. Capocasa,³⁰ F. Carbognani,³⁴ S. Caride,⁷¹ J. Casanueva Diaz,²³ C. Casentini,^{25,13} S. Caudill,¹⁶ M. Cavaglia,²¹ F. Cavalier,²³ R. Cavalieri,³⁴ G. Cella,¹⁹ C. B. Cepeda,¹ L. Cerboni Baiardi,^{57,58} G. Cerretani,^{18,19} E. Cesarini,^{25,13} R. Chakraborty,¹ T. Chalermongsak,¹ S. J. Chamberlin,⁷² M. Chan,³⁶ S. Chao,⁷³ P. Charlton,⁷⁴ E. Chassande-Mottin,³⁰ H. Y. Chen,⁷⁵ Y. Chen,⁷⁶ C. Cheng,⁷³ A. Chincarini,⁴⁷ A. Chiummo,³⁴ H. S. Cho,⁷⁷ M. Cho,⁶² J. H. Chow,²⁰ N. Christensen,^{78,53} Q. Chu,⁵¹ S. Chua,⁶⁰ S. Chung,⁵¹ G. Ciani,⁵ F. Clara,³⁷ J. A. Clark,⁶³ F. Cleva,⁵³ E. Coccia,^{25,12,13} P.-F. Cohadon,⁶⁰ A. Colla,^{79,28} C. G. Collette,⁸⁰ L. Cominsky,⁸¹ M. Constancio Jr.,¹¹ A. Conte,^{79,28} L. Conti,⁴² D. Cook,³⁷ T. R. Corbitt,² N. Cornish,³¹ A. Corsi,⁷¹ S. Cortese,³⁴ C. A. Costa,¹¹ M. W. Coughlin,⁷⁸ S. B. Coughlin,⁸² J.-P. Coulon,⁵³ S. T. Countryman,³⁹ P. Couvares,¹ E. E. Cowan,⁶³ D. M. Coward,⁵¹ M. J. Cowart,⁶ D. C. Coyne,¹ R. Coyne,⁷¹ K. Craig,³⁶ J. D. E. Creighton,¹⁶ J. Cripe,² S. G. Crowder,⁸³ A. Cumming,³⁶ L. Cunningham,³⁶ E. Cuoco,³⁴ T. Dal Canton,⁸ S. L. Danilishin,³⁶ S. D'Antonio,¹³ K. Danzmann,^{17,8} N. S. Darman,⁸⁴ V. Dattilo,³⁴ I. Dave,⁴⁸ H. P. Daveloza,⁸⁵ M. Davier,²³ G. S. Davies,³⁶ E. J. Daw,⁸⁶ R. Day,³⁴ D. DeBra,⁴⁰ G. Debreczeni,³⁸ J. Degallaix,⁶⁵ M. De Laurentis,^{67,4} S. Deléglise,⁶⁰ W. Del Pozzo,⁴⁵ T. Denker,^{8,17} T. Dent,⁸ H. Dereli,⁵³ V. Dergachev,¹ R. T. DeRosa,⁶ R. De Rosa,^{67,4} R. DeSalvo,⁸⁷ S. Dhurandhar,¹⁴ M. C. Díaz,⁸⁵ L. Di Fiore,⁴ M. Di Giovanni,^{79,28} A. Di Lieto,^{18,19} S. Di Pace,^{79,28} I. Di Palma,^{29,8} A. Di Virgilio,¹⁹ G. Dojcinoski,⁸⁸ V. Dolique,⁶⁵ F. Donovan,¹⁰ K. L. Dooley,²¹ S. Doravari,^{6,8} R. Douglas,³⁶ T. P. Downes,¹⁶ M. Drago,^{8,89,90} R. W. P. Drever,¹ J. C. Driggers,³⁷ Z. Du,⁷⁰ M. Ducrot,⁷ S. E. Dwyer,³⁷ T. B. Edo,⁸⁶ M. C. Edwards,⁷⁸ A. Effler,⁶ H.-B. Eggenstein,⁸ P. Ehrens,¹ J. Eichholz,⁵ S. S. Eikenberry,⁵ W. Engels,⁷⁶ R. C. Essick,¹⁰ T. Etzel,¹ M. Evans,¹⁰ T. M. Evans,⁶ R. Everett,⁷² M. Factourovich,³⁹ V. Fafone,^{25,13,12} H. Fair,³⁵ S. Fairhurst,⁹¹ X. Fan,⁷⁰ Q. Fang,⁵¹ S. Farinon,⁴⁷ B. Farr,⁷⁵ W. M. Farr,⁴⁵ M. Favata,⁸⁸ M. Fays,⁹¹ H. Fehrmann,⁸ M. M. Fejer,⁴⁰ I. Ferrante,^{18,19} E. C. Ferreira,¹¹ F. Ferrini,³⁴ F. Fidecaro,^{18,19} I. Fiori,³⁴ D. Fiorucci,³⁰ R. P. Fisher,³⁵ R. Flaminio,^{65,92} M. Fletcher,³⁶ J.-D. Fournier,⁵³ S. Franco,²³ S. Frasca,^{79,28} F. Frasconi,¹⁹ Z. Frei,⁵⁴ A. Freise,⁴⁵ R. Frey,⁵⁹ V. Frey,²³ T. T. Fricke,⁸ P. Fritschel,¹⁰ V. V. Frolov,⁶ P. Fulda,⁵ M. Fyffe,⁶ H. A. G. Gabbard,²¹ J. R. Gair,⁹³ L. Gammaitoni,^{32,33} S. G. Gaonkar,¹⁴ F. Garufi,^{67,4} A. Gatto,³⁰ G. Gaur,^{94,95} N. Gehrels,⁶⁸ G. Gemme,⁴⁷ B. Gendre,⁵³ E. Genin,³⁴ A. Gennai,¹⁹ J. George,⁴⁸ L. Gergely,⁹⁶ V. Germain,⁷ Archisman Ghosh,¹⁵ S. Ghosh,^{52,9} J. A. Giaime,^{2,6} K. D. Giardino,⁶ A. Giazotto,¹⁹ K. Gill,⁹⁷ A. Glaefke,³⁶ E. Goetz,⁹⁸ R. Goetz,⁵ L. Gondan,⁵⁴ G. González,² J. M. Gonzalez Castro,^{18,19} A. Gopakumar,⁹⁹ N. A. Gordon,³⁶ M. L. Gorodetsky,⁴⁹ S. E. Gossan,¹ M. Gosselin,³⁴ R. Gouaty,⁷ C. Graef,³⁶ P. B. Graff,⁶² M. Granata,⁶⁵ A. Grant,³⁶ S. Gras,¹⁰ C. Gray,³⁷ G. Greco,^{57,58} A. C. Green,⁴⁵ P. Groot,⁵² H. Grote,⁸ S. Grunewald,²⁹ G. M. Guidi,^{57,58} X. Guo,⁷⁰ A. Gupta,¹⁴ M. K. Gupta,⁹⁵ K. E. Gushwa,¹ E. K. Gustafson,¹ R. Gustafson,⁹⁸ J. J. Hacker,²² B. R. Hall,⁵⁶ E. D. Hall,¹ G. Hammond,³⁶ M. Haney,⁹⁹ M. M. Hanke,⁸ J. Hanks,³⁷

C. Hanna,⁷² M. D. Hannam,⁹¹ J. Hanson,⁶ T. Hardwick,² J. Harms,^{57,58} G. M. Harry,¹⁰⁰ I. W. Harry,²⁹ M. J. Hart,³⁶ M. T. Hartman,⁵ C.-J. Haster,⁴⁵ K. Haughian,³⁶ A. Heidmann,⁶⁰ M. C. Heintze,^{5,6} H. Heitmann,⁵³ P. Hello,²³ G. Hemming,³⁴ M. Hendry,³⁶ I. S. Heng,³⁶ J. Hennig,³⁶ A. W. Heptonstall,¹ M. Heurs,^{8,17} S. Hild,³⁶ D. Hoak,¹⁰¹ K. A. Hodge,¹ D. Hofman,⁶⁵ S. E. Hollitt,¹⁰² K. Holt,⁶ D. E. Holz,⁷⁵ P. Hopkins,⁹¹ D. J. Hosken,¹⁰² J. Hough,³⁶ E. A. Houston,³⁶ E. J. Howell,⁵¹ Y. M. Hu,³⁶ S. Huang,⁷³ E. A. Huerta,^{103,82} D. Huet,²³ B. Hughey,⁹⁷ S. Husa,⁶⁶ S. H. Huttner,³⁶ T. Huynh-Dinh,⁶ A. Idrisy,⁷² N. Indik,⁸ D. R. Ingram,³⁷ R. Inta,⁷¹ H. N. Isa,³⁶ J.-M. Isac,⁶⁰ M. Isi,¹ G. Islas,²² T. Isogai,¹⁰ B. R. Iyer,¹⁵ K. Izumi,³⁷ T. Jacqmin,⁶⁰ H. Jang,⁷⁷ K. Jani,⁶³ P. Jaranowski,¹⁰⁴ S. Jawahar,¹⁰⁵ F. Jiménez-Forteza,⁶⁶ W. W. Johnson,² D. I. Jones,²⁶ R. Jones,³⁶ R. J. G. Jonker,⁹ L. Ju,⁵¹ Haris K,¹⁰⁶ C. V. Kalaghatgi,^{24,91} V. Kalogera,⁸² S. Kandhasamy,²¹ G. Kang,⁷⁷ J. B. Kanner,¹ S. Karki,⁵⁹ M. Kasprzack,^{2,23,34} E. Katsavounidis,¹⁰ W. Katzman,⁶ S. Kaufer,¹⁷ T. Kaur,⁵¹ K. Kawabe,³⁷ F. Kawazoe,^{8,17} F. Kéfélian,⁵³ M. S. Kehl,⁶⁹ D. Keitel,^{8,66} D. B. Kelley,³⁵ W. Kells,¹ R. Kennedy,⁸⁶ J. S. Key,⁸⁵ A. Khalaidovski,⁸ F. Y. Khalili,⁴⁹ I. Khan,¹² S. Khan,⁹¹ Z. Khan,⁹⁵ E. A. Khazanov,¹⁰⁷ N. Kijbunchoo,³⁷ C. Kim,⁷⁷ J. Kim,¹⁰⁸ K. Kim,¹⁰⁹ Nam-Gyu Kim,⁷⁷ Namjun Kim,⁴⁰ Y.-M. Kim,¹⁰⁸ E. J. King,¹⁰² P. J. King,³⁷ D. L. Kinzel,⁶ J. S. Kissel,³⁷ L. Kleybolte,²⁷ S. Klimenko,⁵ S. M. Koehlenbeck,⁸ K. Kokeyama,² S. Koley,⁹ V. Kondrashov,¹ A. Kontos,¹⁰ M. Korobko,²⁷ W. Z. Korth,¹ I. Kowalska,⁴⁴ D. B. Kozak,¹ V. Kringel,⁸ A. Królak,^{110,111} C. Krueger,¹⁷ G. Kuehn,⁸ P. Kumar,⁶⁹ L. Kuo,⁷³ A. Kutynia,¹¹⁰ B. D. Lackey,³⁵ M. Landry,³⁷ J. Lange,¹¹² B. Lantz,⁴⁰ P. D. Lasky,¹¹³ A. Lazzarini,¹ C. Lazzaro,^{63,42} P. Leaci,^{29,79,28} S. Leavey,³⁶ E. O. Lebigot,^{30,70} C. H. Lee,¹⁰⁸ H. K. Lee,¹⁰⁹ H. M. Lee,¹¹⁴ K. Lee,³⁶ A. Lenon,³⁵ M. Leonardi,^{89,90} J. R. Leong,⁸ N. Leroy,²³ N. Letendre,⁷ Y. Levin,¹¹³ B. M. Levine,³⁷ T. G. F. Li,¹ A. Libson,¹⁰ T. B. Littenberg,¹¹⁵ N. A. Lockerbie,¹⁰⁵ J. Logue,³⁶ A. L. Lombardi,¹⁰¹ J. E. Lord,³⁵ M. Lorenzini,^{12,13} V. Lorette,¹¹⁶ M. Lormand,⁶ G. Losurdo,⁵⁸ J. D. Lough,^{8,17} H. Lück,^{17,8} A. P. Lundgren,⁸ J. Luo,⁷⁸ R. Lynch,¹⁰ Y. Ma,⁵¹ T. MacDonald,⁴⁰ B. Machenschalk,⁸ M. MacInnis,¹⁰ D. M. Macleod,² F. Magaña-Sandoval,³⁵ R. M. Magee,⁵⁶ M. Mageswaran,¹ E. Majorana,²⁸ I. Maksimovic,¹¹⁶ V. Malvezzi,^{25,13} N. Man,⁵³ I. Mandel,⁴⁵ V. Mandic,⁸³ V. Mangano,³⁶ G. L. Mansell,²⁰ M. Manske,¹⁶ M. Mantovani,³⁴ F. Marchesoni,^{117,33} F. Marion,⁷ S. Márka,³⁹ Z. Márka,³⁹ A. S. Markosyan,⁴⁰ E. Maros,¹ F. Martelli,^{57,58} L. Martellini,⁵³ I. W. Martin,³⁶ R. M. Martin,⁵ D. V. Martynov,¹ J. N. Marx,¹ K. Mason,¹⁰ A. Masserot,⁷ T. J. Massinger,³⁵ M. Masso-Reid,³⁶ F. Matichard,¹⁰ L. Matone,³⁹ N. Mavalvala,¹⁰ N. Mazumder,⁵⁶ G. Mazzolo,⁸ R. McCarthy,³⁷ D. E. McClelland,²⁰ S. McCormick,⁶ S. C. McGuire,¹¹⁸ G. McIntyre,¹ J. McIver,¹ D. J. McManus,²⁰ S. T. McWilliams,¹⁰³ D. Meacher,⁷² G. D. Meadors,^{29,8} J. Meidam,⁹ A. Melatos,⁸⁴ G. Mendell,³⁷ D. Mendoza-Gandara,⁸ R. A. Mercer,¹⁶ E. Merilh,³⁷ M. Merzougui,⁵³ S. Meshkov,¹ C. Messenger,³⁶ C. Messick,⁷² P. M. Meyers,⁸³ F. Mezzani,^{28,79} H. Miao,⁴⁵ C. Michel,⁶⁵ H. Middleton,⁴⁵ E. E. Mikhailov,¹¹⁹ L. Milano,^{67,4} J. Miller,¹⁰ M. Millhouse,³¹ Y. Minenkov,¹³ J. Ming,^{29,8} S. Mirshekari,¹²⁰ C. Mishra,¹⁵ S. Mitra,¹⁴ V. P. Mitrofanov,⁴⁹ G. Mitselmakher,⁵ R. Mittleman,¹⁰ A. Moggi,¹⁹ M. Mohan,³⁴ S. R. P. Mohapatra,¹⁰ M. Montani,^{57,58} B. C. Moore,⁸⁸ C. J. Moore,¹²¹ D. Moraru,³⁷ G. Moreno,³⁷ S. R. Morriss,⁸⁵ K. Mossavi,⁸ B. Mours,⁷ C. M. Mow-Lowry,⁴⁵ C. L. Mueller,⁵ G. Mueller,⁵ A. W. Muir,⁹¹ Arunava Mukherjee,¹⁵ D. Mukherjee,¹⁶ S. Mukherjee,⁸⁵ N. Mukund,¹⁴ A. Mullavey,⁶ J. Munch,¹⁰² D. J. Murphy,³⁹ P. G. Murray,³⁶ A. Mytidis,⁵ I. Nardecchia,^{25,13} L. Naticchioni,^{79,28} R. K. Nayak,¹²² V. Necula,⁵ K. Nedkova,¹⁰¹ G. Nelemans,^{52,9} M. Neri,^{46,47} A. Neunzert,⁹⁸ G. Newton,³⁶ T. T. Nguyen,²⁰ A. B. Nielsen,⁸ S. Nissanke,^{52,9} A. Nitz,⁸ F. Nocera,³⁴ D. Nolting,⁶ M. E. N. Normandin,⁸⁵ L. K. Nuttall,³⁵ J. Oberling,³⁷ E. Ochsner,¹⁶ J. O'Dell,¹²³ E. Oelker,¹⁰ G. H. Ogin,¹²⁴ J. J. Oh,¹²⁵ S. H. Oh,¹²⁵ F. Ohme,⁹¹ M. Oliver,⁶⁶ P. Oppermann,⁸ Richard J. Oram,⁶ B. O'Reilly,⁶ R. O'Shaughnessy,¹¹² D. J. Ottaway,¹⁰² R. S. Ottens,⁵ H. Overmier,⁶ B. J. Owen,⁷¹ A. Pai,¹⁰⁶ S. A. Pai,⁴⁸ J. R. Palamos,⁵⁹ O. Palashov,¹⁰⁷ C. Palomba,²⁸ A. Pal-Singh,²⁷ H. Pan,⁷³ C. Pankow,⁸² F. Pannarale,⁹¹ B. C. Pant,⁴⁸ F. Paoletti,^{34,19} A. Paoli,³⁴ M. A. Papa,^{29,16,8} H. R. Paris,⁴⁰ W. Parker,⁶ D. Pascucci,³⁶ A. Pasqualetti,³⁴ R. Passaquieti,^{18,19} D. Passuello,¹⁹ B. Patricelli,^{18,19} Z. Patrick,⁴⁰ B. L. Pearlstone,³⁶ M. Pedraza,¹ R. Pedurand,⁶⁵ L. Pekowsky,³⁵ A. Pele,⁶ S. Penn,¹²⁶ A. Perreca,¹ M. Phelps,³⁶ O. Piccinni,^{79,28} M. Pichot,⁵³ F. Piergiovanni,^{57,58} V. Pierro,⁸⁷ G. Pillant,³⁴ L. Pinard,⁶⁵ I. M. Pinto,⁸⁷ M. Pitkin,³⁶ R. Poggiani,^{18,19} P. Popolizio,³⁴ A. Post,⁸ J. Powell,³⁶ J. Prasad,¹⁴ V. Predoi,⁹¹ S. S. Premachandra,¹¹³ T. Prestegard,⁸³ L. R. Price,¹ M. Prijatelj,³⁴ M. Principe,⁸⁷ S. Privitera,²⁹ G. A. Prodi,^{89,90} L. Prokhorov,⁴⁹ O. Puncken,⁸ M. Punturo,³³ P. Puppo,²⁸ M. Pürerer,²⁹ H. Qi,¹⁶ J. Qin,⁵¹ V. Quetschke,⁸⁵ E. A. Quintero,¹ R. Quitzow-James,⁵⁹ F. J. Raab,³⁷ D. S. Rabeling,²⁰ H. Radkins,³⁷ P. Raffai,⁵⁴ S. Raja,⁴⁸ M. Rakhmanov,⁸⁵ P. Rapagnani,^{79,28} V. Raymond,²⁹ M. Razzano,^{18,19} V. Re,²⁵ J. Read,²² C. M. Reed,³⁷ T. Regimbau,⁵³ L. Rei,⁴⁷ S. Reid,⁵⁰ D. H. Reitze,^{1,5} H. Rew,¹¹⁹ S. D. Reyes,³⁵ F. Ricci,^{79,28} K. Riles,⁹⁸ N. A. Robertson,^{1,36} R. Robie,³⁶ F. Robinet,²³ A. Rocchi,¹³

L. Rolland,⁷ J. G. Rollins,¹ V. J. Roma,⁵⁹ J. D. Romano,⁸⁵ R. Romano,^{3,4} G. Romanov,¹¹⁹ J. H. Romie,⁶ D. Rosińska,^{127,43} S. Rowan,³⁶ A. Rüdiger,⁸ P. Ruggi,³⁴ K. Ryan,³⁷ S. Sachdev,¹ T. Sadecki,³⁷ L. Sadeghian,¹⁶ L. Salconi,³⁴ M. Saleem,¹⁰⁶ F. Salemi,⁸ A. Samajdar,¹²² L. Sammut,^{84,113} E. J. Sanchez,¹ V. Sandberg,³⁷ B. Sandeen,⁸² J. R. Sanders,^{98,35} B. Sassolas,⁶⁵ B. S. Sathyaprakash,⁹¹ P. R. Saulson,³⁵ O. Sauter,⁹⁸ R. L. Savage,³⁷ A. Sawadsky,¹⁷ P. Schale,⁵⁹ R. Schilling[†],⁸ J. Schmidt,⁸ P. Schmidt,^{1,76} R. Schnabel,²⁷ R. M. S. Schofield,⁵⁹ A. Schönbeck,²⁷ E. Schreiber,⁸ D. Schuette,^{8,17} B. F. Schutz,^{91,29} J. Scott,³⁶ S. M. Scott,²⁰ D. Sellers,⁶ D. Sentenac,³⁴ V. Sequino,^{25,13} A. Sergeev,¹⁰⁷ G. Serna,²² Y. Setyawati,^{52,9} A. Sevigny,³⁷ D. A. Shaddock,²⁰ S. Shah,^{52,9} M. S. Shahriar,⁸² M. Shaltev,⁸ Z. Shao,¹ B. Shapiro,⁴⁰ P. Shawhan,⁶² A. Sheperd,¹⁶ D. H. Shoemaker,¹⁰ D. M. Shoemaker,⁶³ K. Siellez,^{53,63} X. Siemens,¹⁶ D. Sigg,³⁷ A. D. Silva,¹¹ D. Simakov,⁸ A. Singer,¹ L. P. Singer,⁶⁸ A. Singh,^{29,8} R. Singh,² A. Singhal,¹² A. M. Sintes,⁶⁶ B. J. J. Slagmolen,²⁰ J. R. Smith,²² N. D. Smith,¹ R. J. E. Smith,¹ E. J. Son,¹²⁵ B. Sorazu,³⁶ F. Sorrentino,⁴⁷ T. Souradeep,¹⁴ A. K. Srivastava,⁹⁵ A. Staley,³⁹ M. Steinke,⁸ J. Steinlechner,³⁶ S. Steinlechner,³⁶ D. Steinmeyer,^{8,17} B. C. Stephens,¹⁶ R. Stone,⁸⁵ K. A. Strain,³⁶ N. Straniero,⁶⁵ G. Stratta,^{57,58} N. A. Strauss,⁷⁸ S. Strigin,⁴⁹ R. Sturani,¹²⁰ A. L. Stuver,⁶ T. Z. Summerscales,¹²⁸ L. Sun,⁸⁴ P. J. Sutton,⁹¹ B. L. Swinkels,³⁴ M. J. Szczepańczyk,⁹⁷ M. Tacca,³⁰ D. Talukder,⁵⁹ D. B. Tanner,⁵ M. Tápai,⁹⁶ S. P. Tarabrin,⁸ A. Taracchini,²⁹ R. Taylor,¹ T. Theeg,⁸ M. P. Thirugnanasambandam,¹ E. G. Thomas,⁴⁵ M. Thomas,⁶ P. Thomas,³⁷ K. A. Thorne,⁶ K. S. Thorne,⁷⁶ E. Thrane,¹¹³ S. Tiwari,¹² V. Tiwari,⁹¹ K. V. Tokmakov,¹⁰⁵ C. Tomlinson,⁸⁶ M. Tonelli,^{18,19} C. V. Torres[‡],⁸⁵ C. I. Torrie,¹ D. Töyrä,⁴⁵ F. Travasso,^{32,33} G. Traylor,⁶ D. Trifirò,²¹ M. C. Tringali,^{89,90} L. Trozzo,^{129,19} M. Tse,¹⁰ M. Turconi,⁵³ D. Tuyenbayev,⁸⁵ D. Ugolini,¹³⁰ C. S. Unnikrishnan,⁹⁹ A. L. Urban,¹⁶ S. A. Usman,³⁵ H. Vahlbruch,¹⁷ G. Vajente,¹ G. Valdes,⁸⁵ N. van Bakel,⁹ M. van Beuzekom,⁹ J. F. J. van den Brand,^{61,9} C. Van Den Broeck,⁹ D. C. Vander-Hyde,^{35,22} L. van der Schaaf,⁹ J. V. van Heijningen,⁹ A. A. van Veggel,³⁶ M. Vardaro,^{41,42} S. Vass,¹ M. Vasúth,³⁸ R. Vaulin,¹⁰ A. Vecchio,⁴⁵ G. Vedovato,⁴² J. Veitch,⁴⁵ P. J. Veitch,¹⁰² K. Venkateswara,¹³¹ D. Verkindt,⁷ F. Vetrano,^{57,58} A. Viceré,^{57,58} S. Vinciguerra,⁴⁵ D. J. Vine,⁵⁰ J.-Y. Vinet,⁵³ S. Vitale,¹⁰ T. Vo,³⁵ H. Vocca,^{32,33} C. Vorvick,³⁷ D. Voss,⁵ W. D. Voudsen,⁴⁵ S. P. Vyatchanin,⁴⁹ A. R. Wade,²⁰ L. E. Wade,¹³² M. Wade,¹³² M. Walker,² L. Wallace,¹ S. Walsh,^{16,8,29} G. Wang,¹² H. Wang,⁴⁵ M. Wang,⁴⁵ X. Wang,⁷⁰ Y. Wang,⁵¹ R. L. Ward,²⁰ J. Warner,³⁷ M. Was,⁷ B. Weaver,³⁷ L.-W. Wei,⁵³ M. Weinert,⁸ A. J. Weinstein,¹ R. Weiss,¹⁰ T. Welborn,⁶ L. Wen,⁵¹ P. Weßels,⁸ T. Westphal,⁸ K. Wette,⁸ J. T. Whelan,^{112,8} D. J. White,⁸⁶ B. F. Whiting,⁵ R. D. Williams,¹ A. R. Williamson,⁹¹ J. L. Willis,¹³³ B. Willke,^{17,8} M. H. Wimmer,^{8,17} W. Winkler,⁸ C. C. Wipf,¹ H. Wittel,^{8,17} G. Woan,³⁶ J. Worden,³⁷ J. L. Wright,³⁶ G. Wu,⁶ J. Yablon,⁸² W. Yam,¹⁰ H. Yamamoto,¹ C. C. Yancey,⁶² M. J. Yap,²⁰ H. Yu,¹⁰ M. Yvert,⁷ A. Zadrożny,¹¹⁰ L. Zangrando,⁴² M. Zanolin,⁹⁷ J.-P. Zendri,⁴² M. Zevin,⁸² F. Zhang,¹⁰ L. Zhang,¹ M. Zhang,¹¹⁹ Y. Zhang,¹¹² C. Zhao,⁵¹ M. Zhou,⁸² Z. Zhou,⁸² X. J. Zhu,⁵¹ M. E. Zucker,^{1,10} S. E. Zuraw,¹⁰¹ and J. Zweizig¹

(LIGO Scientific Collaboration and Virgo Collaboration)

[†]Deceased, May 2015. [‡]Deceased, March 2015.

¹*LIGO, California Institute of Technology, Pasadena, CA 91125, USA*

²*Louisiana State University, Baton Rouge, LA 70803, USA*

³*Università di Salerno, Fisciano, I-84084 Salerno, Italy*

⁴*INFN, Sezione di Napoli, Complesso Universitario di Monte S. Angelo, I-80126 Napoli, Italy*

⁵*University of Florida, Gainesville, FL 32611, USA*

⁶*LIGO Livingston Observatory, Livingston, LA 70754, USA*

⁷*Laboratoire d'Annecy-le-Vieux de Physique des Particules (LAPP),*

Université Savoie Mont Blanc, CNRS/IN2P3, F-74941 Annecy-le-Vieux, France

⁸*Albert-Einstein-Institut, Max-Planck-Institut für Gravitationsphysik, D-30167 Hannover, Germany*

⁹*Nikhef, Science Park, 1098 XG Amsterdam, The Netherlands*

¹⁰*LIGO, Massachusetts Institute of Technology, Cambridge, MA 02139, USA*

¹¹*Instituto Nacional de Pesquisas Espaciais, 12227-010 São José dos Campos, SP, Brazil*

¹²*INFN, Gran Sasso Science Institute, I-67100 L'Aquila, Italy*

¹³*INFN, Sezione di Roma Tor Vergata, I-00133 Roma, Italy*

¹⁴*Inter-University Centre for Astronomy and Astrophysics, Pune 411007, India*

¹⁵*International Centre for Theoretical Sciences, Tata Institute of Fundamental Research, Bangalore 560012, India*

¹⁶*University of Wisconsin-Milwaukee, Milwaukee, WI 53201, USA*

¹⁷*Leibniz Universität Hannover, D-30167 Hannover, Germany*

¹⁸*Università di Pisa, I-56127 Pisa, Italy*

¹⁹*INFN, Sezione di Pisa, I-56127 Pisa, Italy*

²⁰*Australian National University, Canberra, Australian Capital Territory 0200, Australia*

²¹*The University of Mississippi, University, MS 38677, USA*

- ²² *California State University Fullerton, Fullerton, CA 92831, USA*
- ²³ *LAL, Univ. Paris-Sud, CNRS/IN2P3, Université Paris-Saclay, Orsay, France*
- ²⁴ *Chennai Mathematical Institute, Chennai, India*
- ²⁵ *Università di Roma Tor Vergata, I-00133 Roma, Italy*
- ²⁶ *University of Southampton, Southampton SO17 1BJ, United Kingdom*
- ²⁷ *Universität Hamburg, D-22761 Hamburg, Germany*
- ²⁸ *INFN, Sezione di Roma, I-00185 Roma, Italy*
- ²⁹ *Albert-Einstein-Institut, Max-Planck-Institut für Gravitationsphysik, D-14476 Potsdam-Golm, Germany*
- ³⁰ *APC, AstroParticule et Cosmologie, Université Paris Diderot, CNRS/IN2P3, CEA/Irfu, Observatoire de Paris, Sorbonne Paris Cité, F-75205 Paris Cedex 13, France*
- ³¹ *Montana State University, Bozeman, MT 59717, USA*
- ³² *Università di Perugia, I-06123 Perugia, Italy*
- ³³ *INFN, Sezione di Perugia, I-06123 Perugia, Italy*
- ³⁴ *European Gravitational Observatory (EGO), I-56021 Cascina, Pisa, Italy*
- ³⁵ *Syracuse University, Syracuse, NY 13244, USA*
- ³⁶ *SUPA, University of Glasgow, Glasgow G12 8QQ, United Kingdom*
- ³⁷ *LIGO Hanford Observatory, Richland, WA 99352, USA*
- ³⁸ *Wigner RCP, RMKI, H-1121 Budapest, Konkoly Thege Miklós út 29-33, Hungary*
- ³⁹ *Columbia University, New York, NY 10027, USA*
- ⁴⁰ *Stanford University, Stanford, CA 94305, USA*
- ⁴¹ *Università di Padova, Dipartimento di Fisica e Astronomia, I-35131 Padova, Italy*
- ⁴² *INFN, Sezione di Padova, I-35131 Padova, Italy*
- ⁴³ *CAMK-PAN, 00-716 Warsaw, Poland*
- ⁴⁴ *Astronomical Observatory Warsaw University, 00-478 Warsaw, Poland*
- ⁴⁵ *University of Birmingham, Birmingham B15 2TT, United Kingdom*
- ⁴⁶ *Università degli Studi di Genova, I-16146 Genova, Italy*
- ⁴⁷ *INFN, Sezione di Genova, I-16146 Genova, Italy*
- ⁴⁸ *RRCAT, Indore MP 452013, India*
- ⁴⁹ *Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia*
- ⁵⁰ *SUPA, University of the West of Scotland, Paisley PA1 2BE, United Kingdom*
- ⁵¹ *University of Western Australia, Crawley, Western Australia 6009, Australia*
- ⁵² *Department of Astrophysics/IMAPP, Radboud University Nijmegen, P.O. Box 9010, 6500 GL Nijmegen, The Netherlands*
- ⁵³ *Artemis, Université Côte d'Azur, CNRS, Observatoire Côte d'Azur, CS 34229, Nice cedex 4, France*
- ⁵⁴ *MTA Eötvös University, "Lendulet" Astrophysics Research Group, Budapest 1117, Hungary*
- ⁵⁵ *Institut de Physique de Rennes, CNRS, Université de Rennes 1, F-35042 Rennes, France*
- ⁵⁶ *Washington State University, Pullman, WA 99164, USA*
- ⁵⁷ *Università degli Studi di Urbino 'Carlo Bo', I-61029 Urbino, Italy*
- ⁵⁸ *INFN, Sezione di Firenze, I-50019 Sesto Fiorentino, Firenze, Italy*
- ⁵⁹ *University of Oregon, Eugene, OR 97403, USA*
- ⁶⁰ *Laboratoire Kastler Brossel, UPMC-Sorbonne Universités, CNRS, ENS-PSL Research University, Collège de France, F-75005 Paris, France*
- ⁶¹ *VU University Amsterdam, 1081 HV Amsterdam, The Netherlands*
- ⁶² *University of Maryland, College Park, MD 20742, USA*
- ⁶³ *Center for Relativistic Astrophysics and School of Physics, Georgia Institute of Technology, Atlanta, GA 30332, USA*
- ⁶⁴ *Institut Lumière Matière, Université de Lyon, Université Claude Bernard Lyon 1, UMR CNRS 5306, 69622 Villeurbanne, France*
- ⁶⁵ *Laboratoire des Matériaux Avancés (LMA), IN2P3/CNRS, Université de Lyon, F-69622 Villeurbanne, Lyon, France*
- ⁶⁶ *Universitat de les Illes Balears, IAC3—IEEC, E-07122 Palma de Mallorca, Spain*
- ⁶⁷ *Università di Napoli 'Federico II', Complesso Universitario di Monte S. Angelo, I-80126 Napoli, Italy*
- ⁶⁸ *NASA/Goddard Space Flight Center, Greenbelt, MD 20771, USA*
- ⁶⁹ *Canadian Institute for Theoretical Astrophysics, University of Toronto, Toronto, Ontario M5S 3H8, Canada*
- ⁷⁰ *Tsinghua University, Beijing 100084, China*
- ⁷¹ *Texas Tech University, Lubbock, TX 79409, USA*
- ⁷² *The Pennsylvania State University, University Park, PA 16802, USA*
- ⁷³ *National Tsing Hua University, Hsinchu City, Taiwan 30013, R.O.C.*
- ⁷⁴ *Charles Sturt University, Wagga Wagga, New South Wales 2678, Australia*
- ⁷⁵ *University of Chicago, Chicago, IL 60637, USA*
- ⁷⁶ *Caltech CaRT, Pasadena, CA 91125, USA*
- ⁷⁷ *Korea Institute of Science and Technology Information, Daejeon 305-806, Korea*

- ⁷⁸ Carleton College, Northfield, MN 55057, USA
- ⁷⁹ Università di Roma 'La Sapienza', I-00185 Roma, Italy
- ⁸⁰ University of Brussels, Brussels 1050, Belgium
- ⁸¹ Sonoma State University, Rohnert Park, CA 94928, USA
- ⁸² Northwestern University, Evanston, IL 60208, USA
- ⁸³ University of Minnesota, Minneapolis, MN 55455, USA
- ⁸⁴ The University of Melbourne, Parkville, Victoria 3010, Australia
- ⁸⁵ The University of Texas Rio Grande Valley, Brownsville, TX 78520, USA
- ⁸⁶ The University of Sheffield, Sheffield S10 2TN, United Kingdom
- ⁸⁷ University of Sannio at Benevento, I-82100 Benevento, Italy and INFN, Sezione di Napoli, I-80100 Napoli, Italy
- ⁸⁸ Montclair State University, Montclair, NJ 07043, USA
- ⁸⁹ Università di Trento, Dipartimento di Fisica, I-38123 Povo, Trento, Italy
- ⁹⁰ INFN, Trento Institute for Fundamental Physics and Applications, I-38123 Povo, Trento, Italy
- ⁹¹ Cardiff University, Cardiff CF24 3AA, United Kingdom
- ⁹² National Astronomical Observatory of Japan, 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan
- ⁹³ School of Mathematics, University of Edinburgh, Edinburgh EH9 3FD, United Kingdom
- ⁹⁴ Indian Institute of Technology, Gandhinagar Ahmedabad Gujarat 382424, India
- ⁹⁵ Institute for Plasma Research, Bhat, Gandhinagar 382428, India
- ⁹⁶ University of Szeged, Dóm tér 9, Szeged 6720, Hungary
- ⁹⁷ Embry-Riddle Aeronautical University, Prescott, AZ 86301, USA
- ⁹⁸ University of Michigan, Ann Arbor, MI 48109, USA
- ⁹⁹ Tata Institute of Fundamental Research, Mumbai 400005, India
- ¹⁰⁰ American University, Washington, D.C. 20016, USA
- ¹⁰¹ University of Massachusetts-Amherst, Amherst, MA 01003, USA
- ¹⁰² University of Adelaide, Adelaide, South Australia 5005, Australia
- ¹⁰³ West Virginia University, Morgantown, WV 26506, USA
- ¹⁰⁴ University of Białystok, 15-424 Białystok, Poland
- ¹⁰⁵ SUPA, University of Strathclyde, Glasgow G1 1XQ, United Kingdom
- ¹⁰⁶ IISER-TVM, CET Campus, Trivandrum Kerala 695016, India
- ¹⁰⁷ Institute of Applied Physics, Nizhny Novgorod, 603950, Russia
- ¹⁰⁸ Pusan National University, Busan 609-735, Korea
- ¹⁰⁹ Hanyang University, Seoul 133-791, Korea
- ¹¹⁰ NCBJ, 05-400 Świerk-Otwock, Poland
- ¹¹¹ IM-PAN, 00-956 Warsaw, Poland
- ¹¹² Rochester Institute of Technology, Rochester, NY 14623, USA
- ¹¹³ Monash University, Victoria 3800, Australia
- ¹¹⁴ Seoul National University, Seoul 151-742, Korea
- ¹¹⁵ University of Alabama in Huntsville, Huntsville, AL 35899, USA
- ¹¹⁶ ESPCI, CNRS, F-75005 Paris, France
- ¹¹⁷ Università di Camerino, Dipartimento di Fisica, I-62032 Camerino, Italy
- ¹¹⁸ Southern University and A&M College, Baton Rouge, LA 70813, USA
- ¹¹⁹ College of William and Mary, Williamsburg, VA 23187, USA
- ¹²⁰ Instituto de Física Teórica, University Estadual Paulista/ICTP South American Institute for Fundamental Research, São Paulo SP 01140-070, Brazil
- ¹²¹ University of Cambridge, Cambridge CB2 1TN, United Kingdom
- ¹²² IISER-Kolkata, Mohanpur, West Bengal 741252, India
- ¹²³ Rutherford Appleton Laboratory, HSIC, Chilton, Didcot, Oxon OX11 0QX, United Kingdom
- ¹²⁴ Whitman College, 345 Boyer Ave, Walla Walla, WA 99362 USA
- ¹²⁵ National Institute for Mathematical Sciences, Daejeon 305-390, Korea
- ¹²⁶ Hobart and William Smith Colleges, Geneva, NY 14456, USA
- ¹²⁷ Janusz Gil Institute of Astronomy, University of Zielona Góra, 65-265 Zielona Góra, Poland
- ¹²⁸ Andrews University, Berrien Springs, MI 49104, USA
- ¹²⁹ Università di Siena, I-53100 Siena, Italy
- ¹³⁰ Trinity University, San Antonio, TX 78212, USA
- ¹³¹ University of Washington, Seattle, WA 98195, USA
- ¹³² Kenyon College, Gambier, OH 43022, USA
- ¹³³ Abilene Christian University, Abilene, TX 79699, USA ^a

(Dated: February 23, 2016)

The LIGO detection of the gravitational wave transient GW150914, from the inspiral and merger of two black holes with masses $\gtrsim 30 M_{\odot}$, suggests a population of binary black holes with relatively high mass. This observation implies that the stochastic gravitational-wave background from binary black holes, created from the incoherent superposition of all the merging binaries in the Universe,

could be higher than previously expected. Using the properties of GW150914, we estimate the energy density of such a background from binary black holes. In the most sensitive part of the Advanced LIGO/Virgo band for stochastic backgrounds (near 25 Hz), we predict $\Omega_{\text{GW}}(f = 25 \text{ Hz}) = 1.1_{-0.9}^{+2.7} \times 10^{-9}$ with 90% confidence. This prediction is robustly demonstrated for a variety of formation scenarios with different parameters. The differences between models are small compared to the statistical uncertainty arising from the currently poorly constrained local coalescence rate. We conclude that this background is potentially measurable by the Advanced LIGO/Virgo detectors operating at their projected final sensitivity.

PACS numbers: 04.80.Nn, 04.25.dg, 95.85.Sz, 97.80.-d

Introduction — On September 14, 2015 the Advanced LIGO [1, 2] Hanford and Livingston detectors observed the gravitational-wave event GW150914 with a significance in excess of 5.1σ [3]. The observed signal is consistent with a binary black hole waveform with component masses of $m_1 = 36_{-4}^{+5} M_\odot$ and $m_2 = 29_{-4}^{+4} M_\odot$, as measured in the source frame, and coalescing at a luminosity distance of 410_{-180}^{+160} Mpc, corresponding to a redshift of $z = 0.09_{-0.04}^{+0.03}$ [3, 4].

For every event like GW150914 observed by advanced gravitational-wave detectors, there are many more too distant to be resolved. The gravitational waves from these unresolvable events combine to create a stochastic background, which can be detected by correlating the signals from two or more gravitational-wave detectors [5]. While it has long been known that the advanced detectors could observe such a background, the detection of GW150914 suggests that the binary black hole background level is likely to be at the higher end of previous predictions (see, e.g., [6–13]).

Heavy black holes like GW150914 are predicted to form in low-metallicity stellar environments, lower than about half of solar metallicity, and in the presence of relatively weak massive-star winds [14]. These masses are also larger than the masses inferred from reliable dynamical measurements in black-hole X-ray binaries [14]. More massive binaries emit more energy in gravitational waves. Hence, the measurement of the component masses of GW150914 favors a higher amplitude of the corresponding gravitational-wave background.

In addition, the coalescence rate of binary black holes like GW150914 in the local Universe is estimated to be $16_{-13}^{+38} \text{ Gpc}^{-3} \text{ yr}^{-1}$ [15] median with 90% credible interval. This rate excludes the lower end of pre-detection rate estimates [14], while being consistent with the higher end. A higher coalescence rate also implies a brighter stochastic background.

There are currently two possible formation channels that are consistent with the GW150914 event [14]. Binary black holes may be formed from isolated binaries of massive stars in galactic fields, or through dynamical interactions in dense stellar environments such as globular

clusters [14]. The evolution of the merger rate with redshift depends in part on the assumed formation scenario.

In this paper we discuss the detectability of the stochastic background produced by binary black holes throughout the Universe based on the measured properties of GW150914.

Binary black hole background — The energy density spectrum of gravitational waves is described by the following dimensionless quantity [5]:

$$\Omega_{\text{GW}}(f) = \frac{f}{\rho_c} \frac{d\rho_{\text{GW}}}{df}, \quad (1)$$

where $d\rho_{\text{GW}}$ is the energy density in the frequency interval f to $f + df$, $\rho_c = 3H_0^2 c^2 / 8\pi G$ is the critical energy density required to close the Universe, and $H_0 = 67.8 \pm 0.9 \text{ km/s/Mpc}$ is the Hubble constant [16].

A population of binary black holes is characterized by the distribution of the intrinsic source parameters θ (usually the component masses and spin). Since this distribution is unknown at present, following [15] and [17] we divide the distribution into distinct classes corresponding to the observed candidates. If binary black holes in some class k , with source parameters θ_k , merge at a rate $R_m(z; \theta_k)$ per unit comoving volume V_c per unit source time, then the total gravitational-wave energy density spectrum from all the sources in this class is given by (see, e.g. [6–13]):

$$\Omega_{\text{GW}}(f; \theta_k) = \frac{f}{\rho_c H_0} \int_0^{z_{\text{max}}} dz \frac{R_m(z, \theta_k) \frac{dE_{\text{GW}}}{df_s}(f_s, \theta_k)}{(1+z)E(\Omega_M, \Omega_\Lambda, z)}, \quad (2)$$

and the final energy density spectrum is the sum of $\Omega_{\text{GW}}(f; \theta_k)$ from each class.¹ In Eq. 2, $dE_{\text{GW}}/df_s(f_s, \theta_k)$ is the spectral energy density of a source of class k at the frequency $f_s = f(1+z)$, which depends on the source parameters θ_k ; $E(\Omega_M, \Omega_\Lambda, z) = \sqrt{\Omega_M(1+z)^3 + \Omega_\Lambda}$ captures the dependence of the comoving volume on redshift for the standard flat cosmology model, with $\Omega_M = 0.31$

^a spokesperson@ligo.org

¹ When the distribution of the source parameters is better understood after multiple detections, the discrete sum can be replaced by a continuous integral.

and $\Omega_\Lambda = 1 - \Omega_M$. The $(1+z)$ factor in the denominator of Eq. 2 corrects for the cosmic expansion, converting time in the source frame to the detector frame. The parameter z_{\max} corresponds to the time of the first coalescences. We set $z_{\max} = 10$, noting, however, that sources above $z \sim 5$ contribute very little to the total background (see, e.g., [6–13]).

The merger rate $R_m(z; \theta_k)$ is a convolution of the binary formation rate $R_f(z; \theta_k)$ with the distribution of the time delays $P(t_d; \theta_k)$ between binary black hole formation and merger (see e.g., [18])

$$R_m(z; \theta_k) = \int_{t_{\min}}^{t_{\max}} R_f(z_f; \theta_k) P(t_d; \theta_k) dt_d, \quad (3)$$

where t_d is the time delay, z_f is the redshift at the formation time $t_f = t(z) - t_d$, and $t(z)$ is the age of the Universe at merger.

Inference on GW150914 [4], along with expectations that gravitational-wave emission is efficient in circularizing the orbit [14], allows us to restrict our models for dE_{GW}/df_s to circular orbits. Measurements do not constrain the component spins in the orbital plane [4]; we therefore restrict our model to spins (anti-)aligned with the orbital angular momentum, and use the functional form of dE_{GW}/df_s derived in [19]. In addition to the component masses, this model depends on the effective spin parameter along the direction of the orbital angular momentum χ_{eff} , which takes values between -1 (in which both black holes have maximal spins anti-aligned with respect to the orbital angular momentum) and $+1$ (assuming maximally aligned spins) [4].

Fiducial Model — The GW150914 event appears consistent with both the dynamic and field formation channels [14]; however the field channel is currently better described in the stochastic background literature. Thus our *Fiducial* model is inspired by population synthesis studies of field binaries (see [13]).

We assume that the binary black hole formation rate is proportional to the star formation rate (SFR) at metallicity $Z \leq Z_\odot/2$ [14], where Z_\odot is the solar metallicity. That is, to compute the binary black hole formation rate, the SFR is multiplied by the fraction of star formation occurring below the metallicity threshold $Z_c = Z_\odot/2$. For the SFR, we use the recent model [20], referred to here as “Vangioni”, based on the gamma-ray burst rate of [21] and on the normalization described in [22, 23]. We adopt the mean metallicity–redshift relation of [24], rescaled upwards by a factor of 3 to account for local observations [20, 25]. In addition, we assume the metallicity is \log_{10} -normally distributed with a standard deviation of 0.5 around the mean at each redshift [26]. We further assume that the time delay distribution follows $P(t_d) \propto t_d^\alpha$, with $\alpha = -1$ for $t_d > t_{\min}$ [18, 27–33], where $t_{\min} = 50 \text{ Myr}$ is the minimum delay time for a massive

binary to evolve until coalescence [e.g., 34], and a maximum time delay t_{\max} equal to the Hubble time.

The rest of the *Fiducial* model parameters correspond to the median inferred parameters of GW150914: the chirp mass $M_c = 28 M_\odot$, the symmetric mass ratio $\eta \sim 0.25$, and the effective spin parameter $\chi_{\text{eff}} = -0.06$. We normalize the overall merger rate so that the local merger rate at $z = 0$ matches the most conservative median inferred rate, $16 \text{ Gpc}^{-3} \text{ yr}^{-1}$ [15].

Results — We plot $\Omega_{\text{GW}}(f)$ for the *Fiducial* model as a solid blue curve in Fig. 1a. The curve is shown against the pink shaded region, which represents the 90% credible interval statistical uncertainty in the local rate. Considering this uncertainty, we predict $\Omega_{\text{GW}}(f = 25 \text{ Hz}) = 1.1_{-0.9}^{+2.7} \times 10^{-9}$. The spectrum is well approximated by a power law $\Omega_{\text{GW}}(f) \propto f^{2/3}$ at low frequencies where the contribution from the inspiral phase is dominant and the spectral energy density is $dE_{\text{GW}}/df_s = [(G\pi)^{2/3}/3] M_c^{5/3} f_s^{-1/3}$. This power law remains a good approximation until the spectrum reaches a maximum at $f \sim 100 \text{ Hz}$. The shape is in agreement with previous predictions (see, e.g., [7–13]), except that the maximum is shifted to lower frequencies, due to the higher mass considered.

This calculation of $\Omega_{\text{GW}}(f)$ captures the total energy density in gravitational waves generated by binary black hole coalescences. In practice, some of these sources will be individually detected as resolved binaries. We define “the residual background” as the energy density spectrum that excludes potentially resolvable binaries. While the total background is a property of the Universe, the residual background is detector-dependent. As sensitivity improves, the surveyed volume increases, more binaries are resolved and the residual background decreases.

The dashed blue curve in Fig. 1a represents the residual background calculated for the network of the Advanced LIGO [1, 2] and Advanced Virgo [36, 37] detectors at final sensitivity, assuming that a binary black hole signal is detected if it is associated with a single-detector matched filter signal-to-noise ratio of $\rho > 8$ in at least two detectors [38]. The difference between the two curves is about 30% in the sensitive frequency band (10–50 Hz), indicating that the residual background carries complementary information about the binary black hole population. Binaries with the same component masses as GW150914 can be detected at a redshift up to $z \lesssim 1.3$ by advanced detectors operating at design sensitivity if optimally located and oriented (see Fig. 4 of [14]). However, most sources at $z \gtrsim 0.4$ will not be individually resolvable because of unfavorable location and orientation.

The sensitive frequency band of the Advanced LIGO–Virgo network to a gravitational-wave background produced by binary black holes is 10–50 Hz, where $\Omega_{\text{GW}} \sim f^{2/3}$. It corresponds to more than 95% of the accumu-

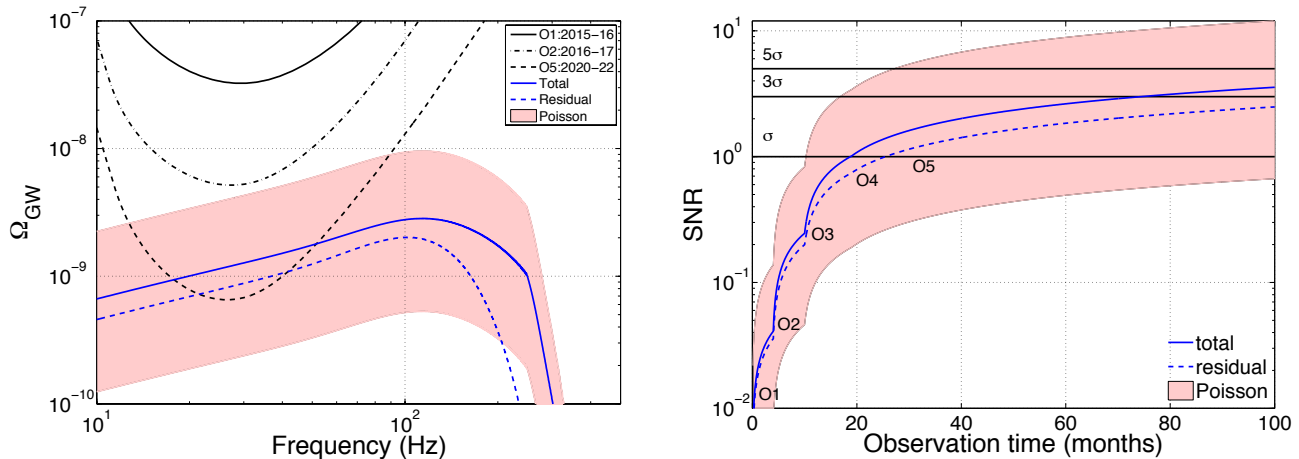


FIG. 1. Expected sensitivity of the network of advanced LIGO and Virgo detectors to the **Fiducial** field model. Left panel: Energy density spectra are shown in blue (solid for the total background; dashed for the residual background, excluding resolved sources, assuming final advanced LIGO and Virgo [1, 2] sensitivity). The pink shaded region “Poisson” shows the 90% CL statistical uncertainty on the total background, propagated from the local rate measurement. The black power-law integrated curves show the 1σ sensitivity of the network expected for the two first observing runs O1 and O2, and for 2 years at the design sensitivity in O5. (O3 and O4 are not significantly different than O5; see Table I.) If the astrophysical background spectrum intersects a black line, it has expected $\text{SNR} \geq 1$. In both panels we assume a coincident duty cycle of 33% for O1 (actual) and 50% for all other runs (predicted). Right panel: Predicted SNR as a function of total observing time. The blue lines and pink shaded region have the same interpretation as in the left panel. Each observing run is indicated by an improvement in the LIGO-Virgo network sensitivity [35], which results in a discontinuity in the slope. The thresholds for $\text{SNR} = 1, 3$ (false-alarm probability $< 3 \times 10^{-3}$) and 5 (false-alarm probability $< 6 \times 10^{-7}$) are indicated by horizontal lines.

lated sensitivity [12, 13, 39]. The black curves shown in Fig. 1a are *power-law integrated curves* [40], which represent the expected 1σ sensitivity of the standard cross-correlation search [5] to power-law gravitational wave backgrounds, of which the $\Omega_{\text{GW}}(f) \propto f^{2/3}$ spectrum for binary inspirals is an example. A power-law integrated curve is calculated by taking the locus of power-law spectra that have expected $\text{SNR} = 1$, where [5]:

$$\text{SNR} = \frac{3H_0^2}{10\pi^2} \sqrt{2T} \left[\int_0^\infty df \sum_{i=1}^n \sum_{j>i} \frac{\gamma_{ij}^2(f) \Omega_{\text{GW}}^2(f)}{f^6 P_i(f) P_j(f)} \right]^{1/2}, \quad (4)$$

for a network of detectors $i = 1, 2, \dots, n$. Hence, if the spectrum of an astrophysical background intersects a black curve, then it has an expected $\text{SNR} \geq 1$. In Eq. 4, $P_i(f)$ and $P_j(f)$ are the one-sided strain noise power spectral densities of two detectors; $\gamma_{ij}(f)$ is the normalized isotropic overlap reduction function [41, 42]; and T is the accumulated coincident observation time. While Eq. 4 is derived by assuming a Gaussian background [5], it can also be applied to non-Gaussian backgrounds (with signals that are clearly separated in time) such as the binary black hole background considered here [43]. The different black curves shown in this plot illustrate the improvement in expected sensitivity in the coming years.

Following [35, 39], we consider five different phases, denoted O1 to O5, corresponding to the first five observing runs, summarized in Table I. For clarity, we show only the O1, O2, and O5 power-law integrated curves since the differences between the projected sensitivities for O3, O4, and O5 are relatively small. In Fig. 1b, we plot the expected accumulated SNR for the **Fiducial** model as a function of total observation time. For both the sensitivity curves and the accumulated SNR, we assume a coincident duty cycle for each pair of detectors of 33% for O1 (actual) and 50% for all other runs (predicted). The total background associated with the **Fiducial** model could be identified with $\text{SNR} = 3$, corresponding to false alarm probability $< 3 \times 10^{-3}$, after approximately 6 years of observing. In the most optimistic scenario given by statistical uncertainties, the total background could be identified after 1.5 years with $\text{SNR} = 3$ and after approximately 2 years with $\text{SNR} = 5$, which is even before design sensitivity is reached. It would take about 2 years of observing to achieve $\text{SNR} = 3$ and about 3.5 years for $\text{SNR} = 5$ for the optimistic residual background. The most pessimistic case considered here is out of reach of the advanced detector network but is in the scope of third generation detectors, such as the proposed Einstein Telescope [?] whose sensitivity would enable to reach $\Omega_{\text{GW}} \sim 10^{-12}$ after a year of observation [? ?].

Alternative Models — We now investigate the impact of

TABLE I. Different phases in the evolution of the aLIGO–AdVirgo detector network over the next several years. The aLIGO and AdVirgo noise curves corresponding to high-sensitivity versions of “Early”, “Mid”, “Late”, and “Design” spectra are taken from [35]. Note that AdVirgo did not participate in the O1 observing run, so is not included in the first phase. “Duration” refers to the planned calendar time as opposed to the amount of accumulated data, for which we assume a duty cycle of 33% for O1 (actual) and 50% for all other runs (predicted). The last column indicates the signal-to-noise ratio at the end of each phase assuming the **Fiducial** model and accounting for the range of uncertainty in the rate.

Observing run	Epoch	Duration (months)	aLIGO sensitivity	AdVirgo sensitivity	SNR (90% CL)
O1	2015–2016	4	Early	—	0–0.14
O2	2016–2017	6	Mid	Early	0.046–0.81
O3	2017–2018	9	Late	Mid	0.19–3.4
O4	2019	12	Design	Late	0.31–5.6
O5	2020+	–	Design	Design	0.67–12

possible variations on the **Fiducial** model. We consider the following alternatives:

- **AltSFR** differs from the **Fiducial** model in assuming a different SFR proposed by Tornatore et al. [44], who combined observations and simulations at higher redshift; the formation rate is assumed to be proportional to the SFR, with no metallicity threshold. We also considered the Madau & Dickinson SFR [24], and found that it produces an energy density spectrum that is essentially indistinguishable from the **Fiducial** model.
- **LongDelay** is identical to the **Fiducial** model but assumes a significantly longer minimum time delay $t_{\min} = 5 \text{ Gyr}$, potentially consistent with binary black hole formation via the chemically homogeneous evolution of rapidly rotating massive stars in very tight binaries [45].
- **LowMetallicity** is the same as **Fiducial**, but assumes that a significantly lower metallicity is required to form heavy black holes, with a threshold of $Z_c = Z_{\odot}/10$ [14].
- **FlatDelay** assumes a flat time delay distribution, $\alpha = 0$, with $t_{\min} = 50 \text{ Myr}$ and $t_{\max} = 1 \text{ Gyr}$. This is inspired by the supposition that dynamical formation of the most massive binaries is likely to happen fairly early in the history of the host environment.
- **ConstRate** follows the assumption of [3] in considering a redshift-independent merger rate, $R_m(z) = 16 \text{ Gpc}^{-3} \text{ yr}^{-1}$.
- **LowMass** is the same as the **Fiducial** model except we add a second class of lower-mass binary black hole sources corresponding to a smaller range for individual detections during O1. As an example, we assume a chirp mass of half the mass of GW150914, $M_c = 15 M_{\odot}$ and a local merger rate of 61 Gpc^{-3}

yr^{-1} [4, 15], corresponding to the second most significant event (LVT151012) identified in [3, 38] with insufficient significance to decisively claim a detection. We assume here that the metallicity threshold is $Z_c = Z_{\odot}$.

Figure 2 shows the impact of alternative models described above. The differences in the spectra of alternative models are not negligible. However, all models considered here fall within the range of statistical uncertainty in the local merger rate estimate relative to the **Fiducial** model in the sensitive frequency band.

The impact of an alternative star formation rate, as examined through model **AltSFR**, is particularly small, indicating that the accuracy of SFR models is not a significant source of systematic error in predicting the strength of the gravitational-wave background.

Relative to the **Fiducial** model, the **LongDelay**, **FlatDelay**, and **ConstRate** models all predict fewer binaries at $z > 0$, even though all of these models are constrained to have the same local merger rate ($z = 0$). These latter three models consequently yield a lower energy density. The **LowMetallicity** model is characterized by a greater high-redshift merger rate than the **Fiducial** model, with significant merger rates extending out to $z \sim 5 - 6$. This is because very little of the local Universe has the required low metallicity, so the local mergers come from the long time-delay tail of a large high-redshift population. Consequently, the **LowMetallicity** model has a higher overall normalization, as well as a different spectral shape at frequencies above 100 Hz due to the redshifting of the dominant high- z contribution to the gravitational-wave background to lower frequencies.

Relative to the **Fiducial** model, the **LowMass** model shows a greater energy density at all frequencies, particularly at high frequencies due to the signals from lower-mass binaries. This model indicates that if there is a significant rate of mergers of binaries with smaller masses than GW150914, their contribution to the gravitational-wave energy density spectrum could be significant. The

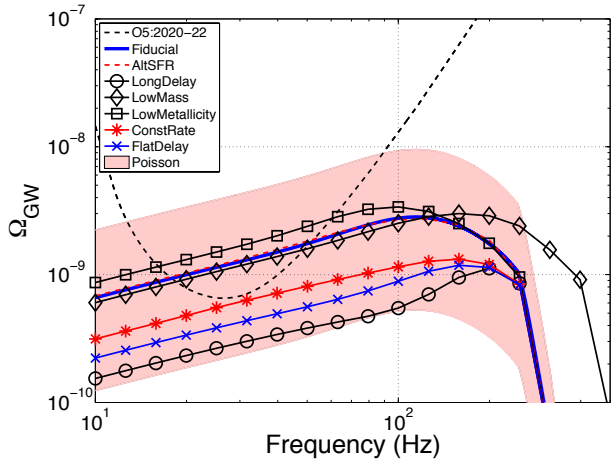


FIG. 2. Energy density spectra for the different models summarized in the text. The pink shaded region “Poisson” shows the 90% CL statistical uncertainty propagated from the local rate measurement, on the **Fiducial** model. The black dashed curve shows the design sensitivity of the network of Advanced LIGO [1, 2] and Virgo [36, 37]; see Tab. I. If the astrophysical background spectrum intersects with the dashed black line, it has expected SNR ≥ 1 .

delta-function mass distributions assumed in all models are motivated by the observed candidates, but are not realistic. We have analyzed two alternative broad mass distributions considered in [15], flat in the log-mass of the component black holes and a Salpeter-like mass function for the larger black hole with a flat mass ratio; these yield broadly consistent energy densities. We have not carried out a systematic study of black hole spin. Measurements of GW150914 prefer small values of spin in the direction of orbital momentum, but spins in the orbital plane are not constrained. Preliminary studies carried out as part of this investigation suggest that $\Omega_{\text{GW}}(f)$ could change by a factor of $\lesssim 2$ for models including spin.

Conclusions and discussion — The detection of gravitational waves from GW150914 is consistent with the existence of high-mass binary black hole mergers with a coalescence rate of tens per Gpc³ per year. As a consequence, the stochastic background from binary black holes is expected to be at the higher end of previous predictions (see, e.g., [7–13]). We have shown that, for the **Fiducial** field model, the energy density spectrum is $\Omega_{\text{GW}}(f = 25 \text{ Hz}) = 1.1^{+2.7}_{-0.9} \times 10^{-9}$ with 90% confidence. This, in turn, implies that the background may be measured by the network of advanced LIGO and Virgo detectors operating at or near their final sensitivity. The uncertainty in this prediction arises from the statistical uncertainty in the local merger rate estimate.

Our predictions are subject to statistical fluctuations in the observed $\Omega_{\text{GW}}(f)$ due to random realizations of the

binaries that coalesce during the observing run. These fluctuations are much smaller than the current local merger uncertainty [43]. The predictions may also be conservative. Throughout, we have assumed the use of the standard cross-correlation statistic, which is known to be sub-optimal for non-Gaussian backgrounds [46]. The development of more sensitive non-Gaussian pipelines may hasten the detection of the binary black hole background [47–49].

We have examined several alternative models for the merger rate evolution with redshift, representative of the uncertainties in the formation channels for high-mass binary black holes. We find that all of these variations lie within the envelope of the uncertain local rate normalization in the 10–50 Hz band, as illustrated in Fig. 2. The power-law slope of the spectrum in this frequency band is not expected to deviate from 2/3 unless there is a significant contribution from sources with high total mass merging at high redshift, $M(1+z) \gtrsim 200M_{\odot}$. This illustrates the robustness of the predicted amplitude and power-law slope of the energy density spectrum.

However, this also implies that the stochastic background measurement with Advanced LIGO and Virgo detectors can only constrain the amplitude of the background power law in the 10–50 Hz sensitive frequency band. The sensitivity of this search at the 2σ level will correspond to $\Omega_{\text{GW}} \sim 10^{-9}$ at 25 Hz with the full-sensitivity network of the Advanced LIGO/Virgo detectors. Therefore, the stochastic search alone will not be able to distinguish between different model variations that have a similar effect on the spectrum in the 10–50 Hz band. Future measurements of individual binary coalescences will help break at least some of these degeneracies, by providing a better estimate of the local merger rate and chirp mass distribution. Combining the two types of measurements (stochastic and individual coalescence event) could therefore help distinguish between different astrophysical formation scenarios for binary black holes [50], but the full potential of this approach may only be reached using third generation of gravitational-wave detectors.

Finally, gravitational waves from distant BBH mergers may be a foreground noise source for the detection of a cosmological background from the early epochs of the Universe in the frequency band of ground-based detectors. However, this astrophysical background has a different spectral shape and different statistical properties (non-continuous and non-Gaussian) that could be used, in principle, to distinguish it from the primordial background.

Acknowledgments — The authors gratefully acknowledge the support of the United States National Science Foundation (NSF) for the construction and operation of the LIGO Laboratory and Advanced LIGO as well as the Science and Technology Facilities Council (STFC) of the United Kingdom, the Max-Planck-Society (MPS),

and the State of Niedersachsen/Germany for support of
the construction of Advanced LIGO and construction
and operation of the GEO600 detector. Additional sup-
port for Advanced LIGO was provided by the Australian
Research Council. The authors gratefully acknowledge
the Italian Istituto Nazionale di Fisica Nucleare (INFN),
the French Centre National de la Recherche Scientifique
(CNRS) and the Foundation for Fundamental Research
on Matter supported by the Netherlands Organisation
for Scientific Research, for the construction and opera-
tion of the Virgo detector and the creation and support
of the EGO consortium. The authors also gratefully ac-
knowledge research support from these agencies as well
as by the Council of Scientific and Industrial Research
of India, Department of Science and Technology, India,
Science & Engineering Research Board (SERB), India,
Ministry of Human Resource Development, India, the
Spanish Ministerio de Economía y Competitividad, the
Conselleria d'Economia i Competitivitat and Conselleria
d'Educació, Cultura i Universitats of the Govern de les
Illes Balears, the National Science Centre of Poland, the
European Commission, the Royal Society, the Scottish
Funding Council, the Scottish Universities Physics Al-
liance, the Hungarian Scientific Research Fund (OTKA),
the Lyon Institute of Origins (LIO), the National Re-
search Foundation of Korea, Industry Canada and the
Province of Ontario through the Ministry of Economic
Development and Innovation, the Natural Science and
Engineering Research Council Canada, Canadian Insti-
tute for Advanced Research, the Brazilian Ministry of
Science, Technology, and Innovation, Russian Founda-
tion for Basic Research, the Leverhulme Trust, the Re-
search Corporation, Ministry of Science and Technology
(MOST), Taiwan and the Kavli Foundation. The au-
thors gratefully acknowledge the support of the NSF,
STFC, MPS, INFN, CNRS and the State of Niedersach-
sen/Germany for provision of computational resources.

[1] G. M. Harry *et al.* (LIGO Scientific Collaboration), Clas-
sical Quantum Gravity **27**, 084006 (2010).
[2] J. Aasi *et al.* (LIGO Scientific Collaboration), Classical
Quantum Gravity **32**, 074001 (2015).
[3] B. Abbott *et al.*, Phys. Rev. Lett. **116**, 061102 (2016),
<https://dcc.ligo.org/LIGO-P150914/public/>.
[4] B. Abbott *et al.*, (2016), <https://dcc.ligo.org/LIGO-P1500218/public/>.
[5] B. Allen and J. D. Romano, Phys. Rev. D **59**, 102001
(1999).
[6] T. Regimbau, Research in Astronomy and Astrophysics
11, 369 (2011).
[7] X.-J. Zhu, E. Howell, T. Regimbau, D. Blair, and Z.-H.
Zhu, Astrophys. J. **739**, 86 (2011).
[8] P. A. Rosado, Phys. Rev. D **84**, 084004 (2011).
[9] S. Marassi, R. Schneider, G. Corvino, V. Ferrari, and
S. P. Zwart, Phys. Rev. D **84**, 124037 (2011).

[10] C. Wu, V. Mandic, and T. Regimbau, Phys. Rev. D **85**,
104024 (2012).
[11] C.-J. Wu, V. Mandic, and T. Regimbau, Phys. Rev. D
87, 042002 (2013).
[12] X.-J. Zhu, E. J. Howell, D. G. Blair, and Z.-H. Zhu,
Monthly Notices of the Royal astronomical Society **431**,
882 (2013).
[13] I. Kowalska-Leszczynska, T. Regimbau, T. Bulik, M. Do-
minik, and K. Belczynski, Astronomy and Astrophysics
574, A58 (2015).
[14] B. Abbott *et al.*, Astrophys. J. Lett. **818** (2016), <https://dcc.ligo.org/LIGO-P1500262/public/>.
[15] B. Abbott *et al.*, (2016), <https://dcc.ligo.org/LIGO-P1500217/public/>.
[16] P. A. R. Ade *et al.* (Planck), (2015), arXiv:1502.01589
[astro-ph.CO].
[17] C. Kim, V. Kalogera, and D. Lorimer, Astrophys. J.
584, 985 (2003).
[18] E. Nakar, Phys. Rep. **442**, 166 (2007).
[19] P. Ajith, Phys. Rev. D **84**, 084037 (2011),
arXiv:1107.1267 [gr-qc].
[20] E. Vangioni, K. A. Olive, T. Prestegard, J. Silk, P. Pe-
titjean, and V. Mandic, Mon. Not. Roy. Astr. Soc. **447**,
2575 (2015).
[21] M. Kistler, H. Yuksel, and A. Hopkins, (2013),
arXiv:1305.1630.
[22] M. Trenti, R. Perna, and S. Tacchella, Astrophys. J.
Supp. **773**, 22 (2013).
[23] P. Behroozi and J. Silk, Astrophys. J. **799**, 32 (2015).
[24] P. Madau and M. Dickinson, AARA **52**, 415 (2014).
[25] Belczynski K., Holz D.E., Bulik T., O'Shaughnessy R.,
(2016), in preparation.
[26] I. Dvorkin, J. Silk, E. Vangioni, P. Petitjean, and
K. A. Olive, Mon. Not. Roy. Astr. Soc. **452**, L36 (2015),
arXiv:1506.06761.
[27] K. Belczynski, V. Kalogera, and T. Bulik, Astrophys. J.
572, 407 (2002).
[28] S. Ando, Journal of Cosmology and Astroparticle Physic
6, 007 (2004).
[29] K. Belczynski, R. Perna, T. Bulik, V. Kalogera,
N. Ivanova, and D. Q. Lamb, Astrophys. J. **648**, 1110
(2006).
[30] J. A. de Freitas Pacheco, T. Regimbau, S. Vincent, and
A. Spallicci, International Journal of Modern Physics D
15, 235 (2006).
[31] E. Berger, D. B. Fox, P. A. Price, E. Nakar, A. Gal-
Yam, D. E. Holz, B. P. Schmidt, A. Cucchiara, S. B.
Cenko, S. R. Kulkarni, A. M. Soderberg, D. A. Frail,
B. E. Penprase, A. Rau, E. Ofek, S. J. B. Burnell, P. B.
Cameron, L. L. Cowie, M. A. Dopita, I. Hook, B. A.
Peterson, P. Podsiadlowski, K. C. Roth, R. E. Rutledge,
S. S. Sheppard, and A. Songaila, Astrophys. J. **664**, 1000
(2007).
[32] R. O'Shaughnessy, K. Belczynski, and V. Kalogera, As-
trophys. J. **675**, 566 (2008).
[33] M. Dominik, K. Belczynski, C. Fryer, D. E. Holz,
E. Berti, T. Bulik, I. Mandel, and R. O'Shaughnessy,
Astrophys. J. **759**, 52 (2012).
[34] M. Dominik, K. Belczynski, C. Fryer, D. E. Holz,
E. Berti, T. Bulik, I. Mandel, and R. O'Shaughnessy,
Astrophys. J. **779**, 72 (2013).
[35] B. P. Abbott *et al.* (LIGO Scientific Collaboration,
Virgo Collaboration), Living Rev. Relat. **19**, 1 (2016),
arXiv:1304.0670 [gr-qc].

- 507 [36] F. Acernese *et al.* (Virgo Collaboration), Classical Quan-521
508 tum Gravity **32**, 024001 (2015). 522
- 509 [37] F. Acernese *et al.* (Virgo Collaboration), Journal of523
510 Physics: Conference Series **610**, 012014 (2015). 524
- 511 [38] B. Abbott *et al.*, (2016), <https://dcc.ligo.org/525>
512 LIGO-P1500269/public/. 526
- 513 [39] D. Meacher, M. Coughlin, S. Morris, T. Regimbau,527
514 N. Christensen, S. Kandhasamy, V. Mandic, J. D. Ro-528
515 mano, and E. Thrane, Phys. Rev. D **92**, 063002 (2015).529
- 516 [40] E. Thrane and J. D. Romano, Phys. Rev. D **88**, 124032530
517 (2013). 531
- 518 [41] E. E. Flanagan, Phys. Rev. D **48**, 2389 (1993). 532
- 519 [42] N. Christensen, Phys. Rev. D **46**, 5250 (1992). 533
- 520 [43] D. Meacher, E. Thrane, and T. Regimbau, Phys. Rev.534
D **89**, 084063 (2014).
- [44] L. Tornatore, S. Borgani, K. Dolag, and F. Matteucci,
Mon. Not. Roy. Astr. Soc. **382**, 1050 (2007).
- [45] I. Mandel and S. E. de Mink, ArXiv e-prints (2016),
arXiv:1601.00007 [astro-ph.HE].
- [46] S. Drasco and E. E. Flanagan, Phys. Rev. D **67**, 082003
(2003).
- [47] N. J. Cornish and J. D. Romano, Phys. Rev. D **92**, 042001
(2015).
- [48] L. Martellini and T. Regimbau, Phys. Rev. D **89**, 124009
(2014).
- [49] E. Thrane, Phys. Rev. D **87**, 043009 (2013).
- [50] V. Mandic, E. Thrane, S. Giampanis, and T. Regimbau,
Phys. Rev. Lett. **109**, 171102 (2012).